

# Estimation of curvature radius of equatorial magnetic field line and its application in M-I mapping

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# Why is the magnetic field-line curvature important

In the nightside tail, the magnetic field-line curvature is (***arguably the most important***) one of the controlling parameters that determine the pitch-angle scattering rate of energetic ions in the equatorial magnetosphere. [Sergeev et al., 1983].

Common threshold of strong/weak pitch angle scattering:

$$\frac{R_c}{\rho_i} > 8 \quad \text{Weak scattering}$$

$$\frac{R_c}{\rho_i} < 8 \quad \text{Strong scattering}$$

# Methods to estimate $R_c$

$$R_c \approx \frac{B_z}{\partial B_x / \partial z} \text{ at the equator}$$

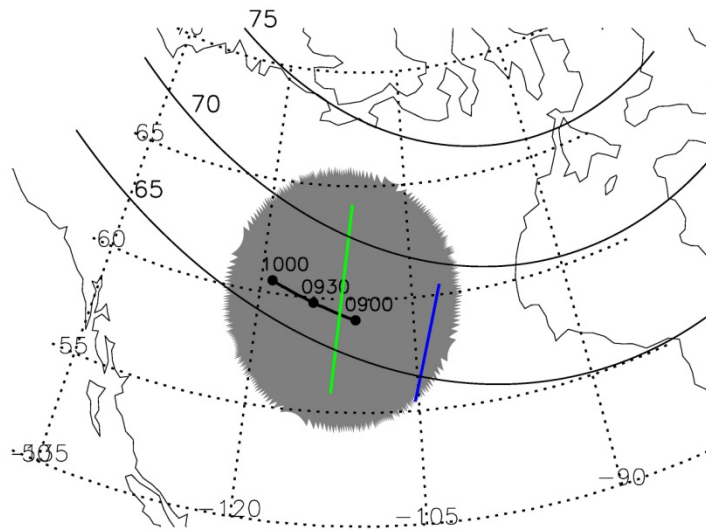
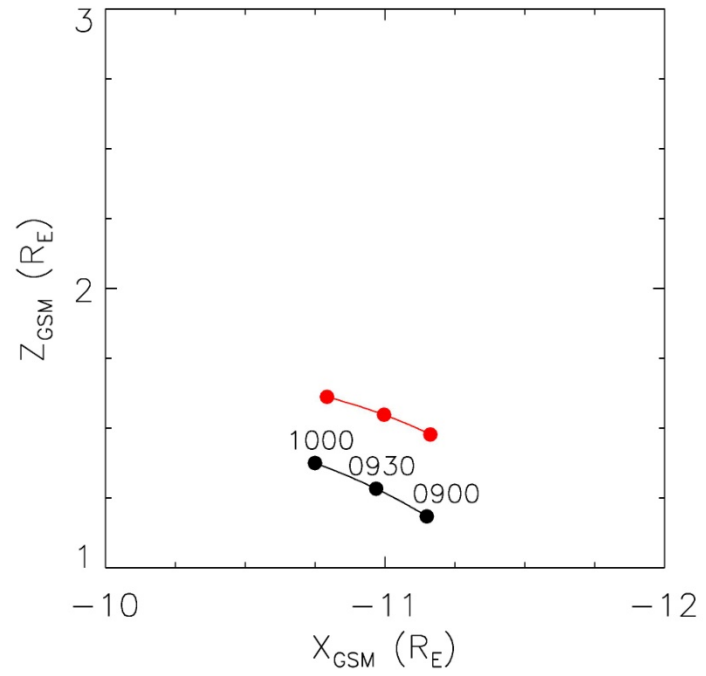
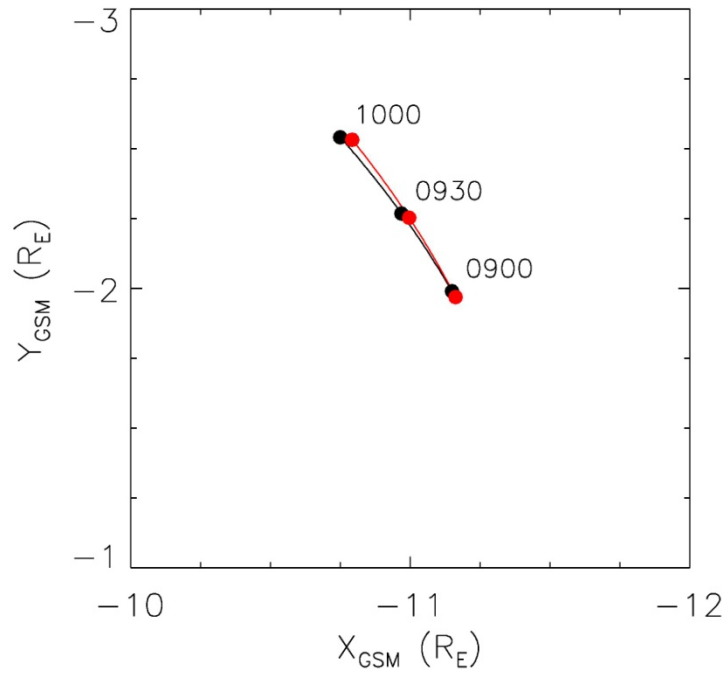
## 1. Single-satellite estimation [Sergeev et al., 1983]

Assumption: 
$$\frac{\partial B_x}{\partial z} \approx \frac{\partial B_x}{\partial t} \cdot \frac{1}{v_z}$$

## 2. Multi-satellite estimation [Donovan et al., 2012; Liang et al., 2012]

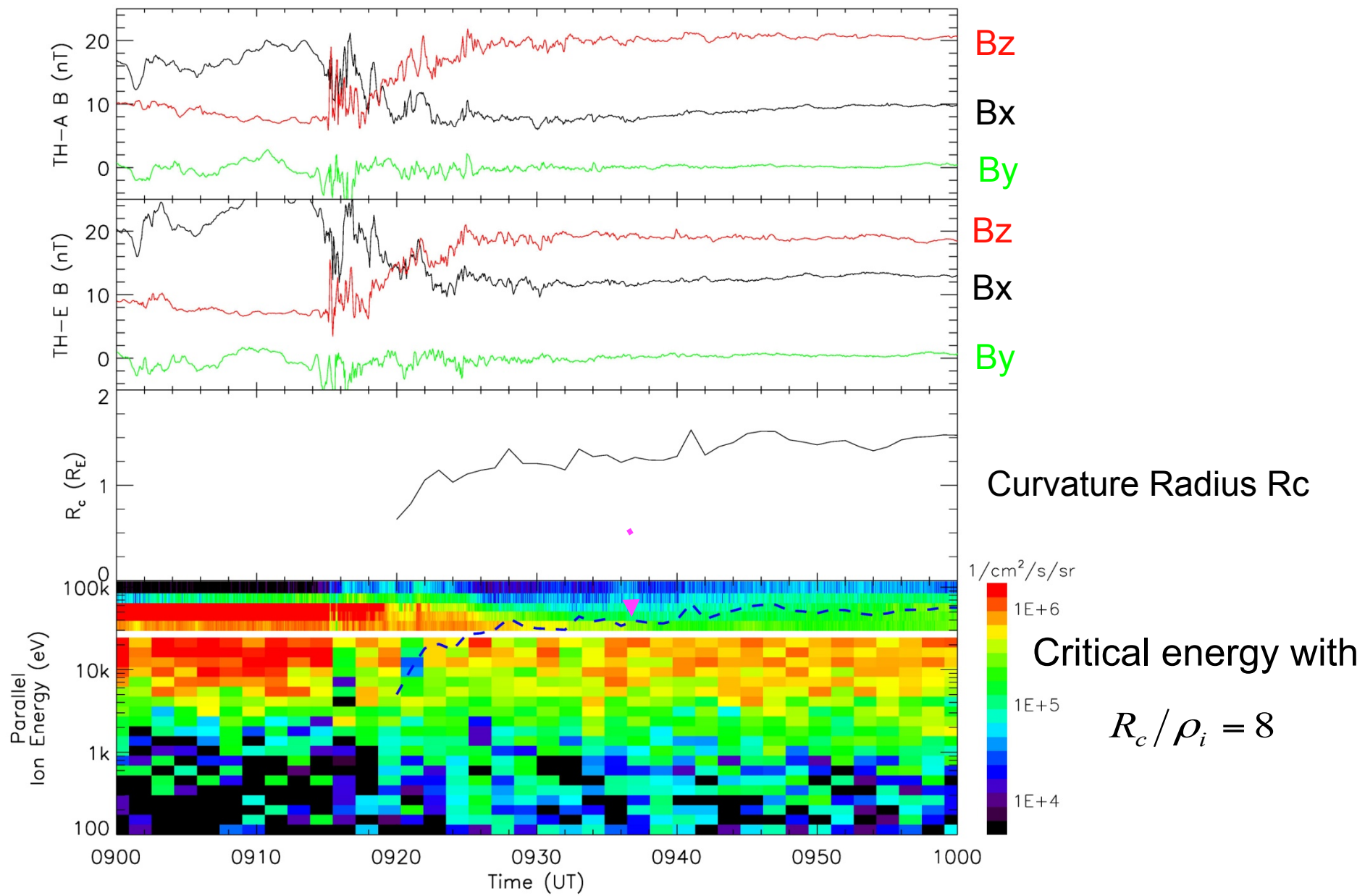
### Requirement and assumption

$$L_c \geq \Delta Z \gg \Delta X \quad \frac{\partial B_x}{\partial z} \gg \left| \frac{\partial B_x}{\partial x} \right| \sim \left| \frac{\partial B_z}{\partial z} \right|$$



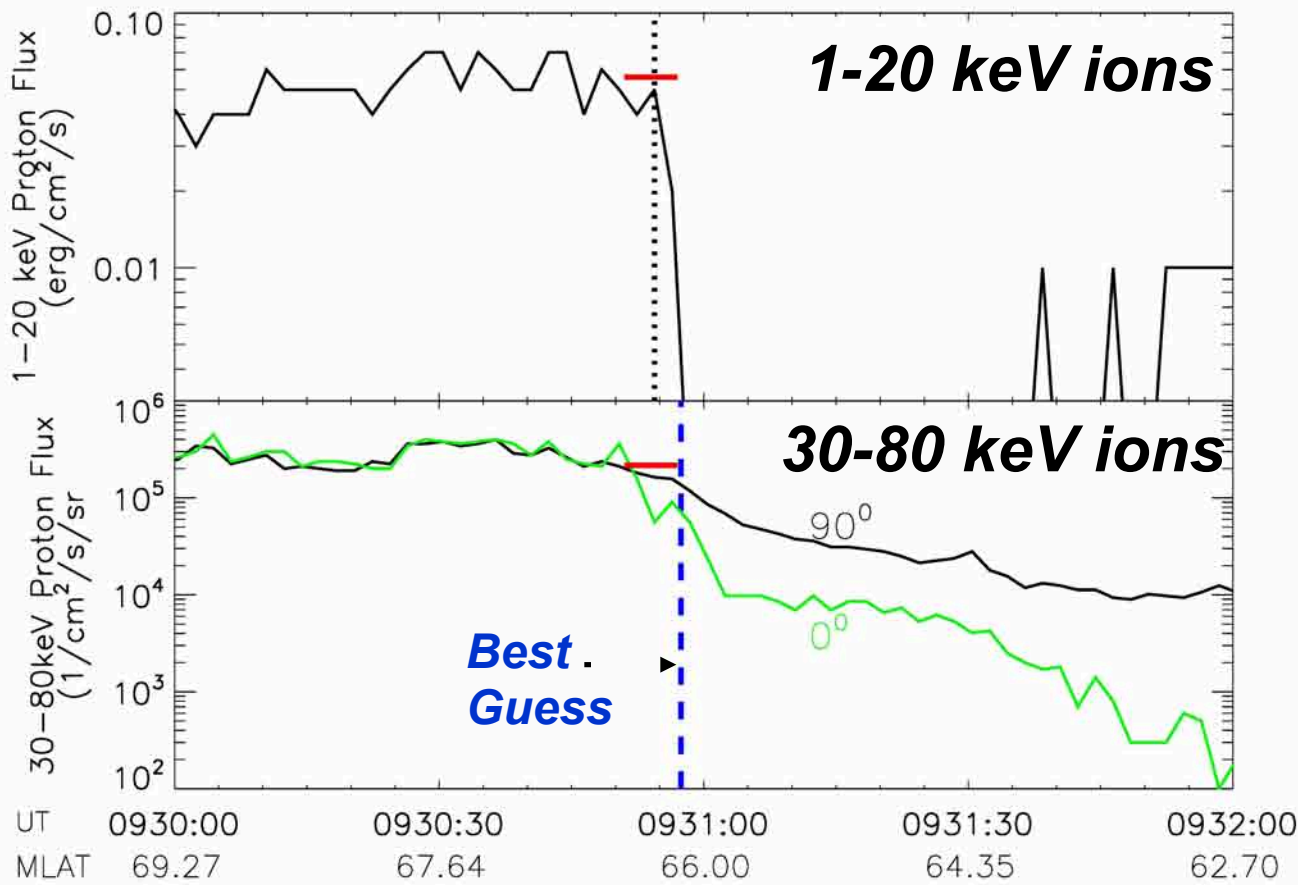
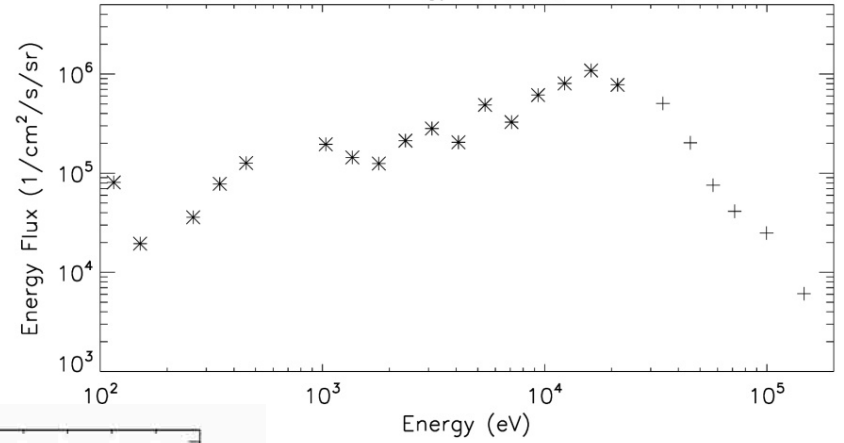
TH-A and E are dominantly separated in Z-distance (~1600 km).

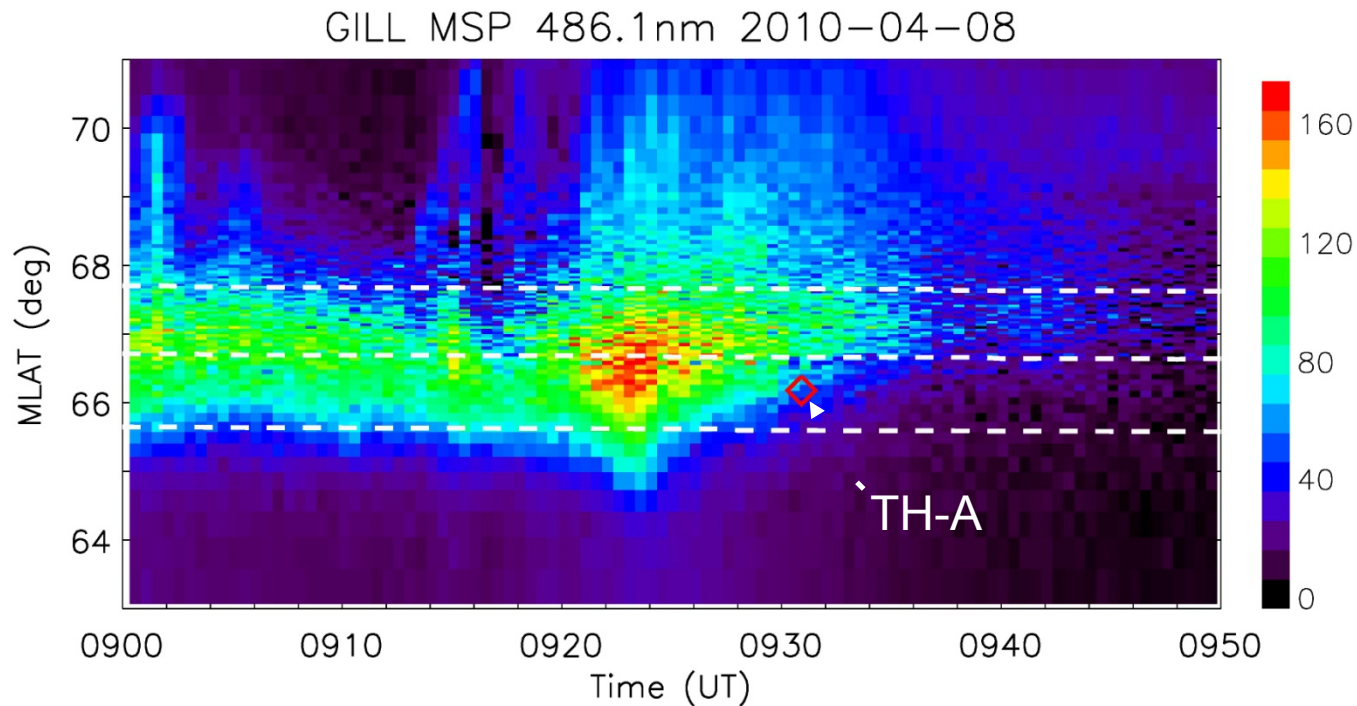
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# NOAA 19 observations and comparison with in-situ parallel fluxes.

TH-A ESA&SST Ion Energy Flux Spectrum at ~0930:48





## Summary:

1. Using the in-situ calculated  $R_c$  one may deduce the PA scattering rate of ions in different energy ranges.
2. Together with the LEO particle and proton aurora measurements, the above information can be used to constrain the mapping of the in-situ probe, and to evaluate the accuracy of mapping based upon empirical field models.